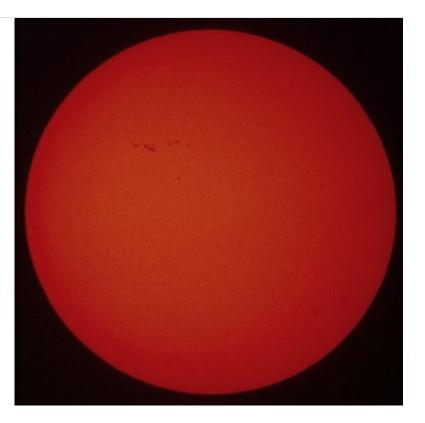
ASTR 1030– October 15

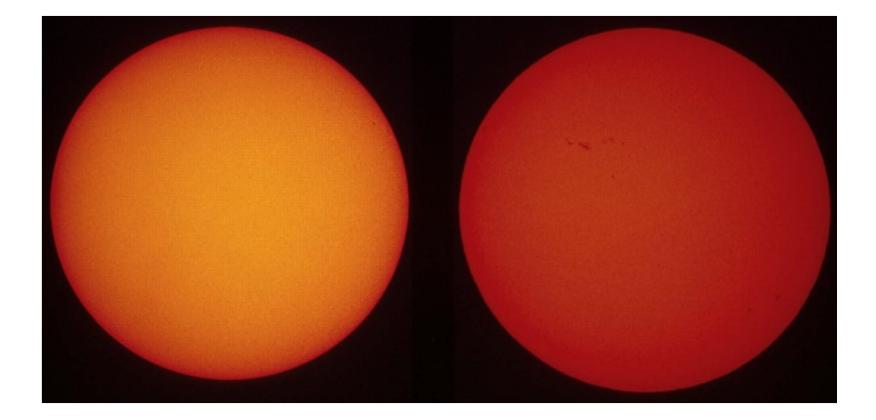
Announcements

Problem Set 4 now posted. Please turn in on Monday October 18 If that's not enough time we'll take it on the 20th

Second Exam on Monday 25th Next Observatory also 25th



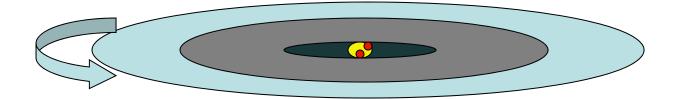
The Sun



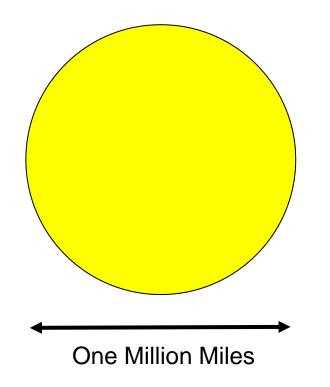
The Sun

Falls into "Disk Stability" 99.9% Ended in Sun (0.1% in Jupiter) Probably the Same Around All Stars -- Planets are Common

> Shoots Planet-Size Bullets into Space



Most of Mass Forms Ball in Center



A Star Is a Hot Ball of Hydrogen (plus 11% Helium)

What Stops the Fall?

Gravity Gets Stronger As Material Gets More Dense

$$F_g = \frac{GMm}{R^2}$$

R smaller implies F greater

The smaller it gets, the faster it falls in!

Why doesn't it just become a black hole?

Or worse yet, a point-like singularity of mass?

Temperature

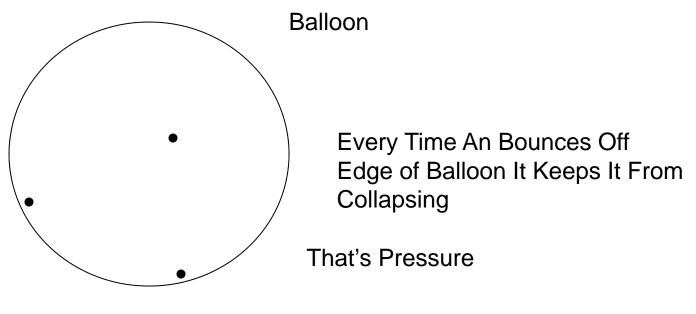
Temperature is a Measure of the Random Kinetic Energy per Particle

The faster the atoms move, the higher the temperature.

But we're talking about *random* motion. If they all move together, then the object moves.

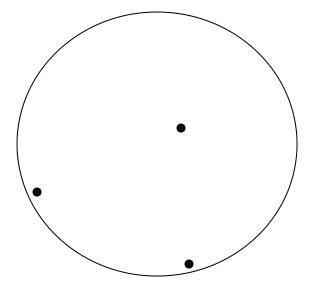
Thermal Pressure

Thermal Gas Pressure Balances Gravitational Pressure



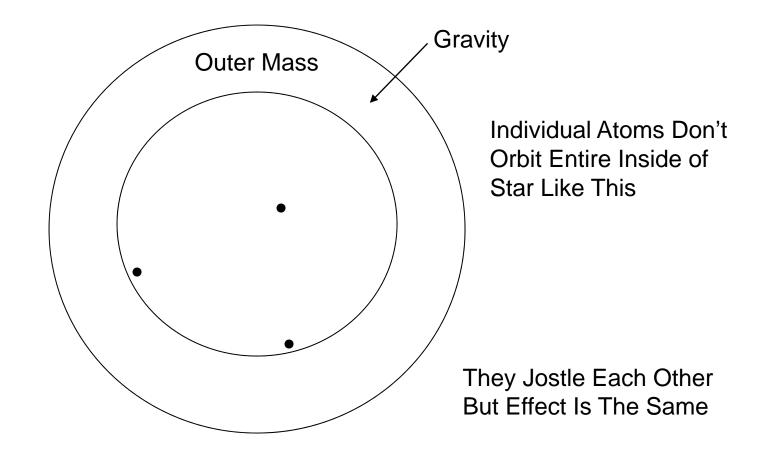
Pressure is Force per Unit Area

Pressure is Proportional to Temperature



Low Temperature Atoms Move Slowly High Temperature Atoms Move Fast

A Star Is Held Up By Thermal Pressure From Below



Temperature Scales

- Fahrenheit
 - 0=salt water freeze 100=human body
- Celsius
 - 0=pure water freeze 100=water boil (sea level)
 - C=(F-32)x5/9
- Kelvin
 - 0=absolute zero 100 degrees between freeze and boil
 - K=C+273
 - -273C = 0K = Absolute Zero

At Absolute Zero Atoms Stop Moving

Thermal Pressure

PV = nRT

Ideal Gas Law Chemistry Style

- P Pressure
- V Volume
- n # moles
- R Constant
- T Temperature (K)

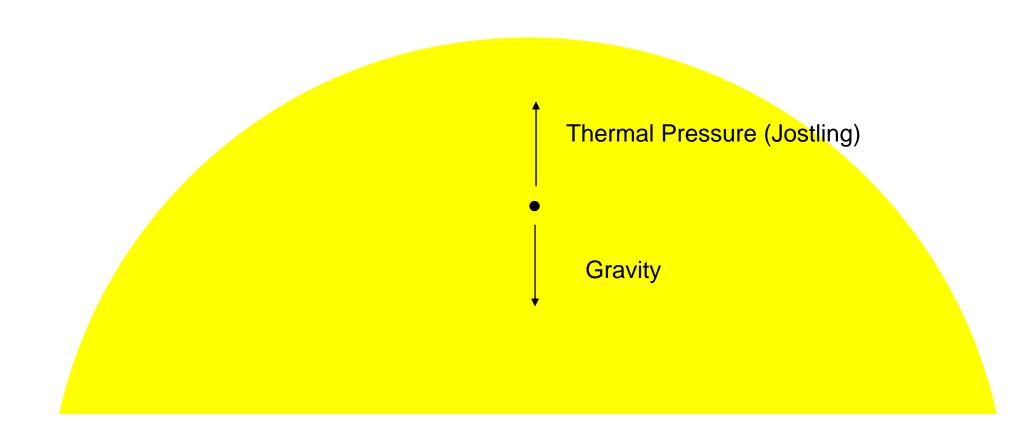
PV = NkT

Ideal Gas Law Physics Style P Pressure
V Volume
N # atoms
k Constant
T Temperature (K)

Pressure Is Proportional to Temperature x Density

Pressure Balance

A Star Always Balances Gravitational Pressure with Thermal Pressure At Each Point Inside



But We Have a Problem

The Sun is Luminous

Radiates Energy Into Space

Luminosity is Power Radiated -- ergs/second

The Energy Comes From Motion of the Atoms Temperature Drops

What Happens When T Drops?

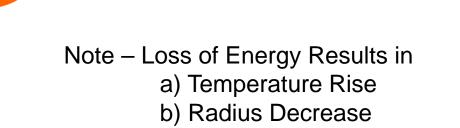
Luminosity Effect

When T Drops Thermal Pressure Can't Hold Off Gravity

The Sun Shrinks -- Radius Drops

Energy is Released as Gas Falls Deeper Into Gravity Field

Temperature Rises



But Wait A Minute... Isn't the Sun Stable?

The Sun has been remarkably stable for 4 billion years as evidenced by geological records.

This collapse is the process by which the Sun coalesced. But then it stopped. Why?

The Sun collapsed until a new source of energy offset the losses to radiation.

NUCLEAR FUSION --- IT'S BURNING HYDROGEN

As long as it burns H at this rate, it will be stable.

Fusion Increases with T

As T in core of Sun increases so does energy production

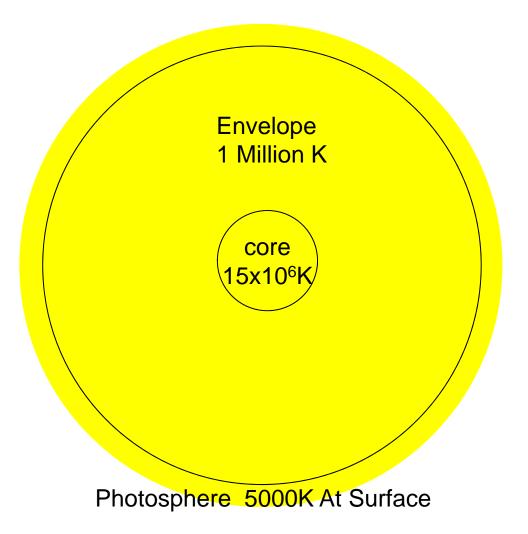
Sun shrank steadily, with T rising until, about 10 million years after it started to form, it reached its current size

There is a VERY fast increase in nuclear energy production above 1,000,000K.

At 15,000,000K in the core nuclear power generated finally balanced the luminosity from the surface.

That's the equilibrium we are still in.

The Nuclear Core



Cosmic Composition

• H	hydrogen	89% by number
• He	helium	11%
• ()	oxygen	0.1%
• C	carbon	0.06%
• N	nitrogen	0.015%

Pretty much the composition of the entire universe. Sun and Jupiter have this composition Earth does not.

Fusion vs. Fission

Fusion: Atoms unite and release energy (Fuse) New atom must be no heavier than iron z=26

ᠲ

Fission: Heavy atoms split to release energy Initial atom must be heavier than iron



WWII Nukes were fission bombs made of U and Pu

Sun works on FUSION of H into He

Proton-Proton Chain

Bottom Line: $H+H+H+H \rightarrow He$

 $_{1}H^{1} + _{1}H^{1} \rightarrow _{1}H^{2} + e^{+} + v$

$$_{1}\text{H}^{2} +_{1}\text{H}^{1} \rightarrow _{2}\text{H}e^{3} + \gamma$$

 $_{2}\text{He}^{3} +_{2}\text{He}^{3} \rightarrow _{2}\text{He}^{4} + _{1}\text{H}^{1} + _{1}\text{H}^{1}$

 $5x10^6 < T < 2x10^7 K$

CNO Cycle

Net: $_{1}H^{1} + _{1}H^{1} + _{1}H^{1} + _{1}H^{1} \rightarrow _{2}He^{4} + 2e + 4\gamma + 2\nu$

hydrogen -> helium + energy

Triple- α Reaction

$$_{2}$$
He⁴ +₂He⁴ \Leftrightarrow $_{4}$ Be⁸ + γ
 $_{4}$ Be⁸ +₂He⁴ \rightarrow $_{6}$ C¹² + γ

 $T < 10^{8}$ K

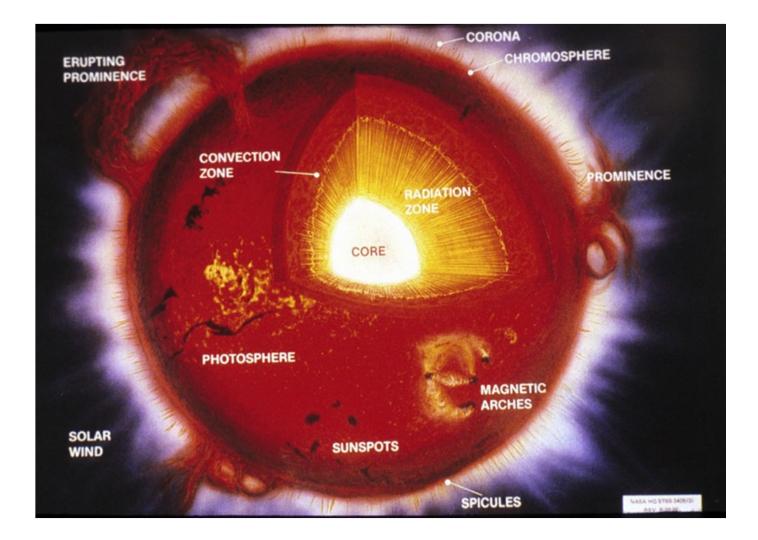
Must be very dense for this to work Be⁸ decays back into helium very quickly unless struck by another He⁴

Too low density in Big Bang

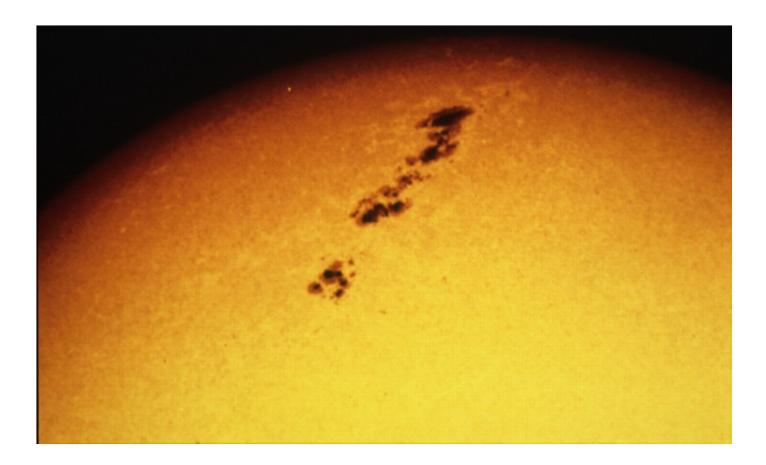
Net:
$$_{2}\text{He}^{4} + _{2}\text{He}^{4} + _{2}\text{He}^{4} + _{2}\text{He}^{4} \rightarrow _{6}\text{C}^{12} + 2\gamma$$

helium -> carbon + energy

Solar Schematic

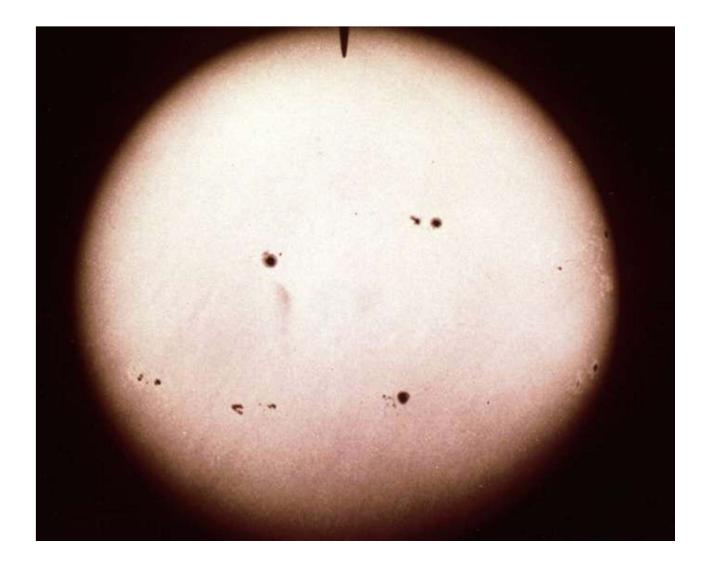


Sunspots



Seen by Ancient Persians

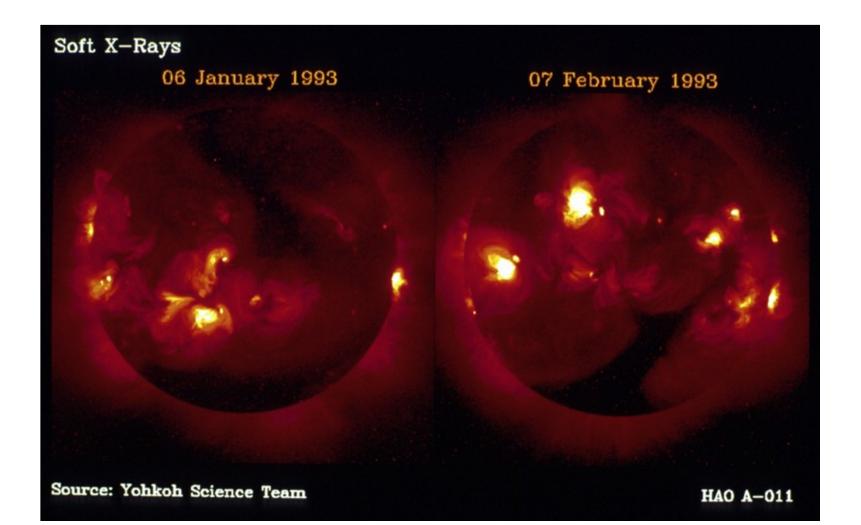
Groups of Sunspots



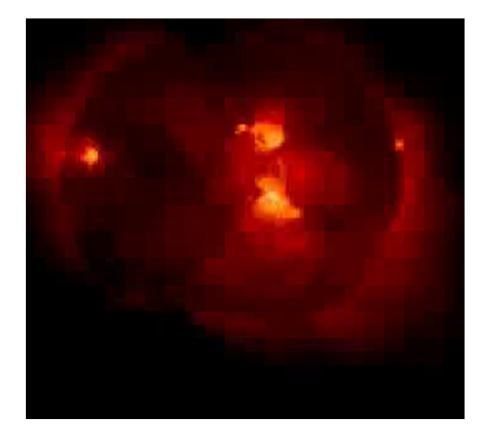
Solar Corona Visible in Eclipse



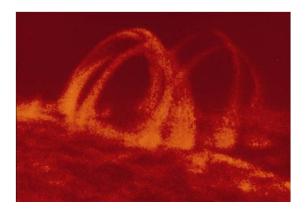
The Sun Viewed in X-rays

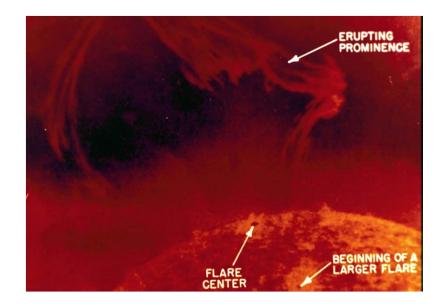


X-ray Movie

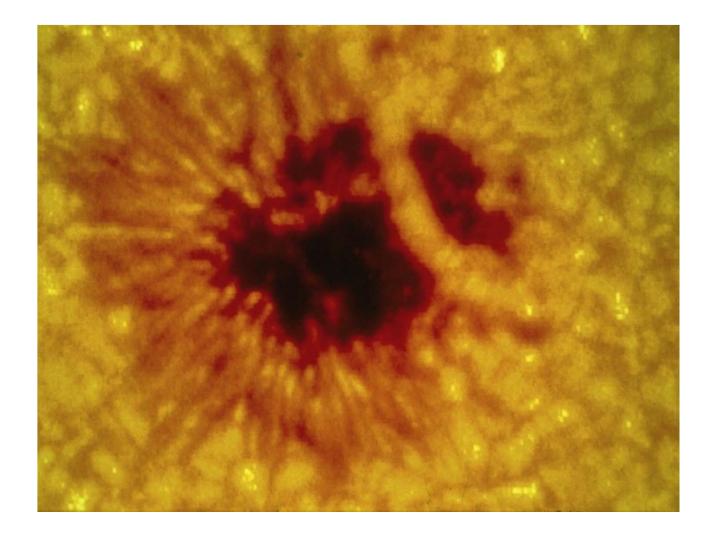


X-ray Loops

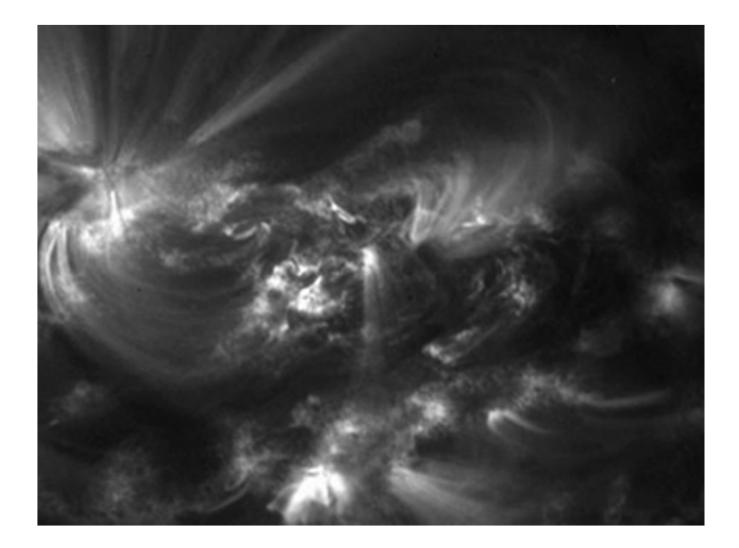


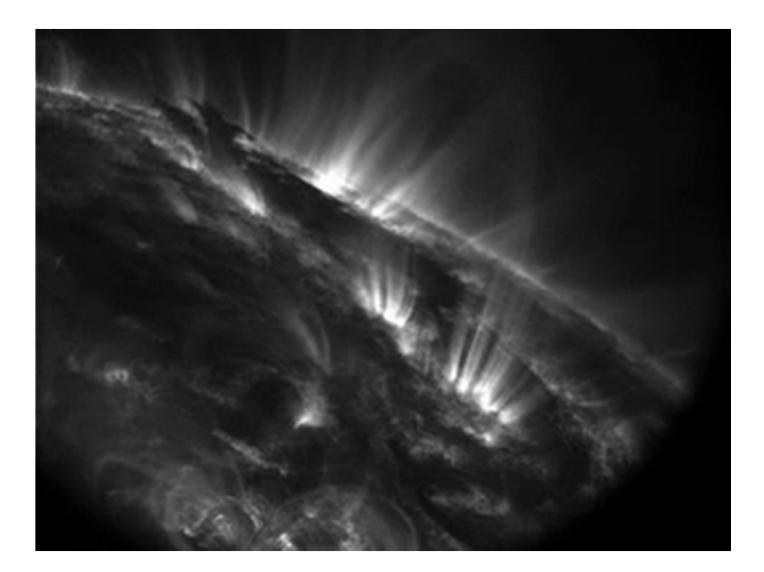


Magnetic Structure

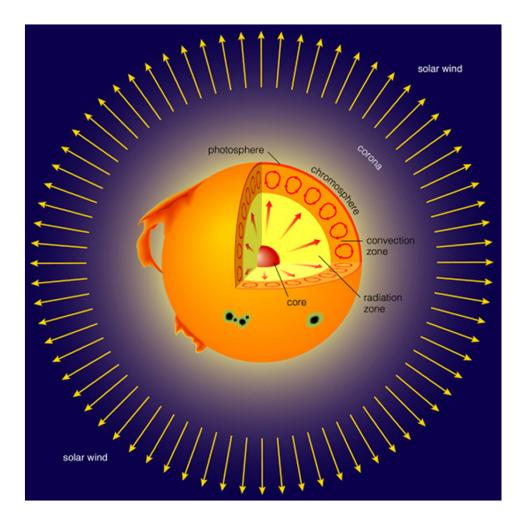


Dynamic Structure



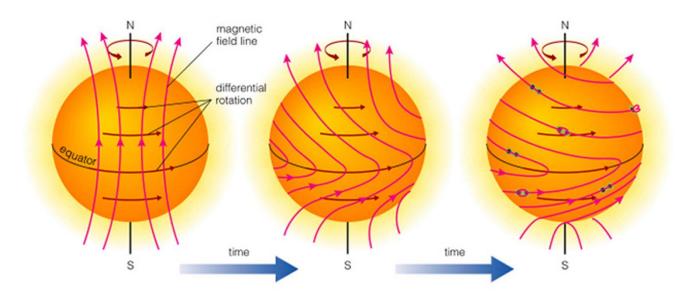


Solar Turbulence

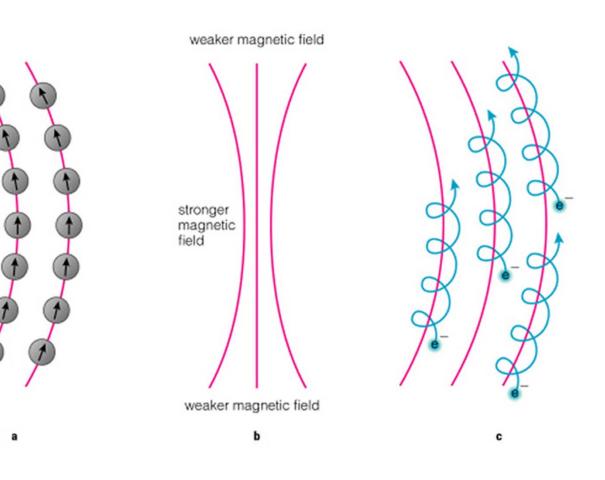


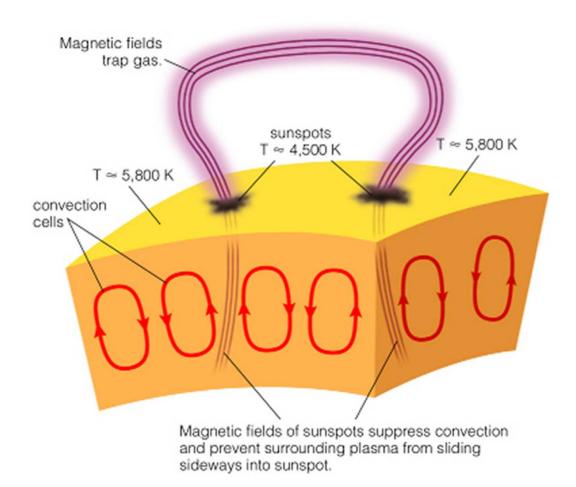
Differential Rotation

Rotates in 25 days at Equator 28 days Mid Latitude 30 days Poles

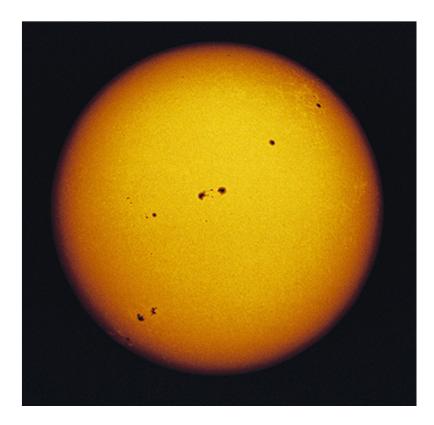


Rapidly Twists Up

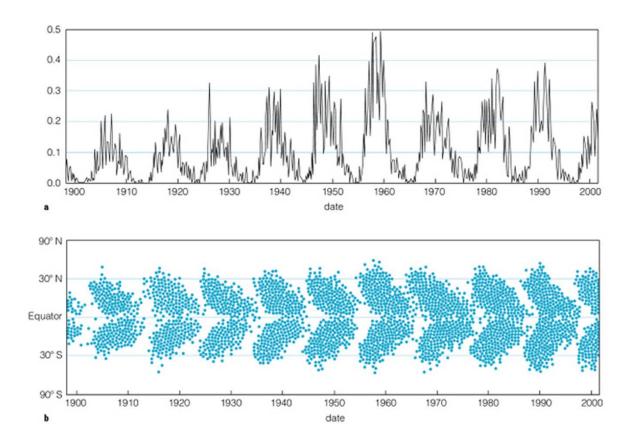




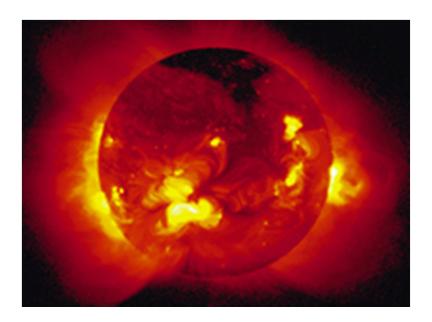
Sunspots Erupt in Groups



Sunspot Cycle



During mid 1600's sunspots became non-existent Maunder Minimum



Solar Wind 5x10⁵K

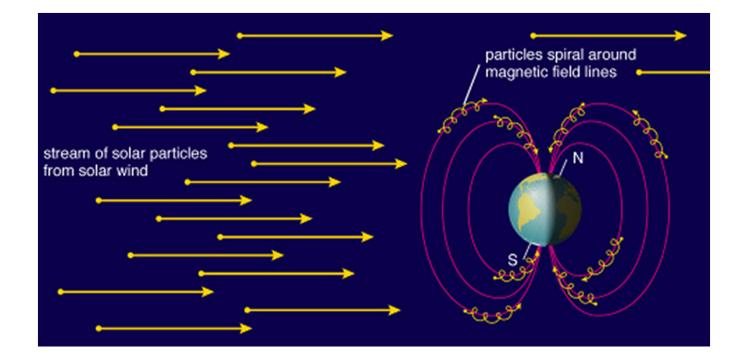
Corona 2x10⁶K

Transition Region 10⁵K

Chromosphere 10⁴K

Photosphere 5500K

Solar Wind Passes Earth



Summary: Sun as a Star

- Formed from cloud 4.6x10⁹ years ago
- Collapsed to present size
 - stabilized by nuclear reactions
- Emits 4x10²⁶W
- Runs on proton-proton chain and CNO cycle
- Now 20% brighter than when first formed
- Turbulent upper envelope
- Magnetic Fields from Differential Rotation
- Sunspots, Corona, Solar Wind
- Activity Cycle 11 years